

Forming the first planetary systems: debris around Galactic thick disc stars

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ABSTRACT

The thick disc contains stars formed within the first Gyr of Galactic history, and little is known about their planetary systems. The *Spitzer* MIPS instrument was used to search 11 of the closest of these old low-metal stars for circumstellar debris, as a signpost that bodies at least as large as planetesimals were formed. A total of 22 thick disc stars has now been observed, after including archival data, but dust is not found in any of the systems. The data rule out a high incidence of debris among star systems from early in the Galaxy's formation. However, some stars of this very old population do host giant planets, at possibly more than the general incidence among low-metal Sun-like stars. As the Solar system contains gas giants but little cometary dust, the thick disc could host analogue systems that formed many Gyr before the Sun.

Key words: planets and satellites: formation – circumstellar matter – infrared: stars.

1 INTRODUCTION

Debris discs are generated by collisions among comets or asteroids, and so are a signpost of planetary systems, indicating that bodies of at least planetesimal size formed. The debris absorbs starlight and re-emits it at longer wavelengths, and observations since 2003 with the *Spitzer Space Telescope* have yielded many new debris detections in the mid- to far-infrared. The relation between giant planets and planetesimal belts is still uncertain. Bryden et al. (2009) find that the incidence and brightness of cool outer-system debris is not affected by the presence or absence of gas giants in the system, which is not surprising if the formation and evolution of giant planets at a few au and comets at tens of au are decoupled processes. More subtle effects may occur, such as the clearing out of comet belts when giant planets migrate, analogous to the Late Heavy Bombardment (LHB) era in Solar system history. This could result in brief dusty episodes as more planetesimals collide, and subsequent lower dust levels due to the ejection of many comets from the system. However, such LHB events seem to be atypical, as *Spitzer* has commonly found old star systems that are still very dusty (Booth et al. 2009).

Greaves et al. (2007) argued that the incidence of giant planet and debris systems, and the weak overlap between them, could be explained if the mass of refractory elements at the planet-formation

stage controlled the outcome. Young circumstellar discs with a large ‘solid’ mass could quickly form planet cores that go on to accrete gas atmospheres, while less substantial discs only form planetesimals by the time the gas disperses and so evolve into debris systems. This model (see also Wyatt, Clarke & Greaves 2007) addresses the differing metallicity dependences observed – metal-rich discs tend to have a higher solid mass and so a greater probability for building a planet core, while lower solid reservoirs are needed to build comets and so debris has little dependence on metallicity (Beichman et al. 2006; Bryden et al. 2006; Greaves et al. 2007; Trilling et al. 2008).

Here we investigate the formation of planetary systems by observing some of the oldest Galactic stars with *Spitzer*. A good test of models is to observe old, metal-poor stars. If these proceeded only as far as planetesimal growth, a high incidence of debris might be seen, but perhaps modified by the longer time available for collisional grinding and dust removal compared to mid-Galactic-age stars like the Sun (Wyatt 2008). The sample of stars studied here is identified as nearby members of the thick disc – stars that formed within about the first Galactic Gyr (Chiappini, Matteucci & Romano 2001) and have orbits inclined with respect to the Galactic plane. These systems are about twice as old as the Solar system.

2 SAMPLE SELECTION

The targets (Table 1) were initially derived from Karataş, Bilir & Schuster (2005), who identified 22 Galactic thick disc stars near the

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Table 1. Properties of the thick disc stars. The first two sets of stars form the volume-limited sample and the third set includes those outside the complete volume or more recently identified as thick disc objects (see text). Metallicities are for Fe except HD 19034 and 101259, which are combined metals. The last columns are the observed date and number of exposures for the new data (see text).

HD	[Fe/H]	d (pc)	Type	Observation date	N_{exp} (24, 70)
New observations					
19034	−0.41	35	G5	27/08/07	2, 20
22879	−0.84	24	F9V	03/03/07	2, 5
29587	−0.61	28	G2V	29/10/07	2, 10
30649	−0.50	30	G1V-VI	29/10/07	2, 15
65583	−0.67	17	G8V	07/11/06	2, 5
68017	−0.42	22	G4V	07/11/06	2, 5
88725	−0.70	36	G1V	29/05/07	2, 10
111515	−0.52	33	G8V	12/07/07	2, 20
144579	−0.70	14	G8V	13/04/07	2, 5
218209	−0.43	30	G6V	16/08/06	2, 20
221830	−0.45	32	F9V	21/08/07	2, 10
Other volume-limited objects					
3795	−0.59	29	K0V		
63077	−0.83	15	F9V		
114729	−0.25	35	G0V		
114762	−0.70	41	F9V		
Other thick disc stars					
6582	−0.86	8	G5Vb		
23439	−1.03	24	K2V, K3V		
76932	−0.86	21	F9V		
101259	−0.23	65	G8IV		
106516	−0.74	23	F5V		
157214	−0.41	14	G0V		

Sun. The identification uses criteria of low metallicity and rotation speed, with values centred around $[\text{Fe}/\text{H}] = -0.5$ and 180 km s^{-1} , respectively, removing halo stars that have lower v_{rot} for the same metallicity. The sample was then reduced to the 15 stars that lie within 40 pc of the Sun, to make the best use of the sensitivity of the MIPS instrument (Rieke et al. 2004). Four of these stars had been observed in previous *Spitzer* programs (Beichman et al. 2006; Trilling et al. 2008; Bryden et al. 2009), and the other 11 were observed for our *Spitzer* program number 30339. Since the survey was proposed, Reddy, Lambert & Allende Prieto (2006) identified further thick disc stars, several of which were observed in various *Spitzer* programs, as was one of the more distant stars of Karataş et al. (2005), HD 101259. The total number of stars discussed here is 22, of which 19 have data at two *Spitzer* wavelengths.

The (logarithmic) metallicities in this sample are -0.23 to -1.03 , considerably lower than the mean of -0.06 with standard deviation of 0.25 for local thin disc stars (Valenti & Fischer 2005). For nine of the volume-limited objects, Takeda et al. (2007) have calculated the ages based on isochrone fitting, a method that is most accurate for old stars that have evolved off the main sequence. Where the age is well determined (uncertainties of $\lesssim 30$ per cent for HD 3795, 30649, 68017, 114729, 218209 and 221830), these stars are dated at 9.6–11.6 Gyr old. This is in very good agreement with their thick disc membership (Karataş et al. 2005), and independent age estimates are consistent with this conclusion. We checked for chromospheric activity, and although current calibrations of age versus activity (e.g. Mamajek & Hillenbrand 2008) are only for roughly solar metallicity, all of our program stars are inferred to have low activity consistent with ages greater than 4 Gyr.

Table 2. Observed fluxes F and errors, and excesses E , R derived from observed fluxes compared to the photospheres. E_{24} is expressed in magnitudes while R_{70} is the ratio of the observed and photospheric fluxes, with a significance for any dust of χ_{70} , i.e. excess flux above the photosphere divided by its error.

HD	F24 (mJy)	F70 (mJy)	E24 (mag)	R70	χ_{70}
3795	137.7 ± 0.7	19.5 ± 2.7	0.03	1.3	1.5
6582	350 ± 1		0.05		
19034	19.9 ± 0.5	0.1 ± 2.4	−0.02	0.0	−0.9
22879	59 ± 0.4	15.4 ± 5.2	−0.04	2.4	1.7
23439A	25.3 ± 0.5		−0.01		
23439B	19.9 ± 0.5		−0.05		
29587	38.8 ± 1.2	2.2 ± 3.9	−0.02	0.5	−0.5
30649	48.2 ± 1.2	1 ± 3.6	0.01	0.2	−1.2
63077	227.8 ± 1.1	15.5 ± 4.1	0.00	0.6	−2.3
65583	65 ± 0.3	12.2 ± 6.5	−0.01	1.7	0.8
68017	64.8 ± 0.3	6.2 ± 6.1	−0.02	0.9	−0.2
76932	133.8 ± 0.7	14.7 ± 2.1	0.03	1.0	−0.0
88725	24.3 ± 0.4	2.3 ± 3.2	−0.03	0.9	−0.1
101259	152.3 ± 0.7	23.3 ± 2.2	0.06	1.4	2.6
106516	80.4 ± 0.4	6.2 ± 6.2	−0.04	0.7	−0.4
111515	20.2 ± 0.5	0.9 ± 2.3	−0.02	0.4	−0.6
114729	62.1 ± 0.7	7.9 ± 2.2	−0.02	1.2	0.5
114762	32.4 ± 0.3	0.2 ± 3.1	−0.05	0.1	−1.1
144579	88.6 ± 0.5	13.1 ± 4.9	−0.02	1.3	0.7
157214	217.5 ± 1	23.9 ± 4.2	0.04	1.0	−0.0
218209	33.4 ± 1.4	0.8 ± 4.2	−0.03	0.2	−0.7
221830	53.5 ± 0.4	9.6 ± 2.7	−0.02	1.6	1.4

3 OBSERVATIONS AND DATA REDUCTION

The 11 new stars were observed between 2006 August and 2007 October, using MIPS at 24 and 70 μm (Table 2). Exposure times were 3 and 10 s at these wavelengths, respectively, multiplied by the number of repeats given in Table 1. The data reduction and instrument calibration procedures are described in Engelbracht et al. (2007) and Gordon et al. (2007). The initial data reduction was done by aperture photometry as described in Trilling et al. (2008). Subsequently, these observations plus the archival data for the other stars were re-reduced homogeneously as part of a *Spitzer* legacy catalogue (Su et al. 2010) and these values are presented.

To identify candidate excesses, photospheric signals were predicted systematically. K -band magnitudes from Two Micron All Sky Survey (2MASS) were extrapolated to 24 μm based on a set of photospheric solar-type stars. A few stars have K -band magnitudes $\lesssim 4.5$ that would saturate the 2MASS images, and for these heritage photometry was transformed into the 2MASS system. These fluxes, for HD 3795, 6582, 63077, 76932 and 157214, will be discussed in the legacy papers (Su, Rieke & Whitelock, in preparation); the last value is rather uncertain with an rms error of ~ 0.06 mag. As no 24- μm excesses were found, the excess at 70 μm was then calculated relative to the 24- μm flux. The significance of excess at 70 μm includes 5 per cent photometric errors plus the statistical error.

One system, HD23439, is a binary with two resolved components that are treated separately here. Amongst the rest of the systems, HD 221830 has a binary companion at approximately 9-arcsec separation that could add 10–15 per cent extra photospheric flux at 70 μm (based on the relative K -band brightnesses), and HD 111515 has a similarly separated companion of a fainter nature, but both of these contaminants are neglected here.

4 RESULTS

4.1 Survey limits

From Table 2, we find no dust excess towards any of the stars. At 24 μm , the differences between observed and photospheric fluxes vary between -0.05 and 0.06 magnitudes, so there are no positive outliers. The standard deviation of the excesses is 0.03 mag and the mean is -0.01 mag. At 70 μm , there is one candidate excess of $+2.6\sigma$, but the most negative value is at -2.3σ , so there are again no notably positive outliers.

About half the stars were not detected at all at 70 μm . This is mainly attributable to larger distances of rare thick disc objects compared to thin disc stars. For example, the median target distance is 28 pc compared to 20 pc in Trilling et al. (2008), while the average 70- μm noise level is about 5 mJy in both cases. Three of the stars (HD 29587, 30649 and 218209) observed here also had high sky backgrounds that can increase the noise. Fig. 1 assesses the sensitivity achieved within the thick disc sample. Survey stars are binned here according to the minimum excess that could be detected with 3σ confidence. The fraction of thick disc stars where only excesses above three times the photosphere can be detected (i.e. bins from 301 to 350 per cent upwards) is rather large. In the general stellar surveys, the sensitivity attained was generally better than this (light blue bins). However, many excesses above 300 per cent were seen in those studies (dark blue bins), suggesting that our survey could be adequately deep if the thick disc stars have typical debris. The biases in comparing the incidence of debris for the thin and thick disc stars are discussed further below.

It was noted that some stars have 70- μm fluxes well *below* the photospheric prediction, although with large errors. Since the MIPS calibration is reliable over a wide range of fluxes (Gordon et al. 2007), we considered whether the stars could be underbright for reasons related to metallicity. However, for a wide sample of stars observed at 24 μm , there is a mild trend of *enhanced* flux at low $[\text{Fe}/\text{H}]$, contributing about 0.02 mag at our median metallicity of -0.6 (Fig. 2). If the 24- μm fluxes are normal or slightly high, it would be surprising if metal-poor stars switched to being under-

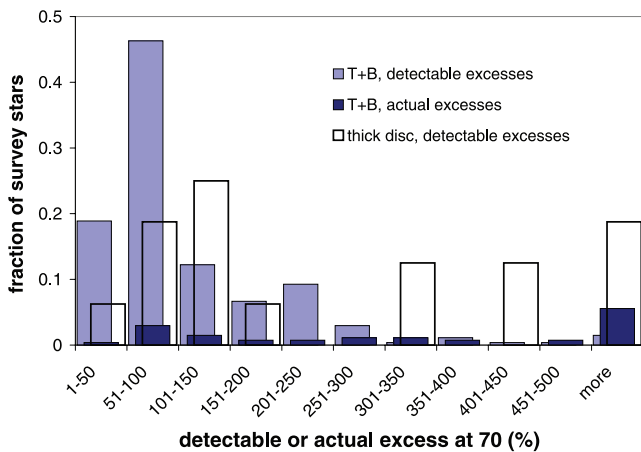


Figure 1. Plot of survey sensitivity at 70 μm (Table 2) compared to Sun-like stars observed by Trilling et al. (2008) and Beichman et al. (2006), denoted by T+B. The bins show relative numbers of survey stars where, if dust was present at that excess level, it could be detected with 3σ confidence. The actual distribution of T+B dust levels is also plotted (dark blue bars). The excesses are above the photosphere, i.e. $R_{70} - 1$.

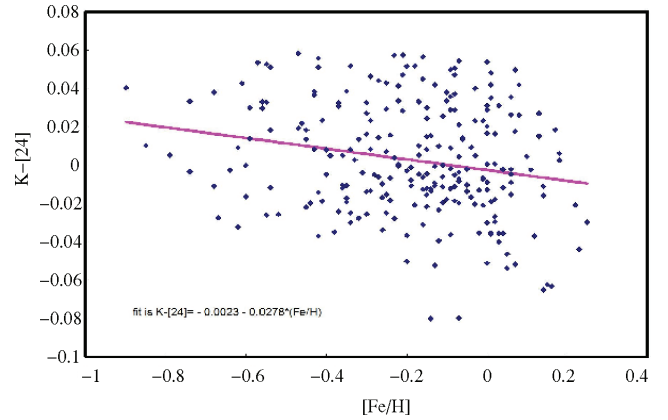


Figure 2. K-[24] colour against $[\text{Fe}/\text{H}]$ (from NSTARS, at <http://nstars.nau.edu>) and linear fit. Data are from various *Spitzer* surveys, re-reduced homogeneously.

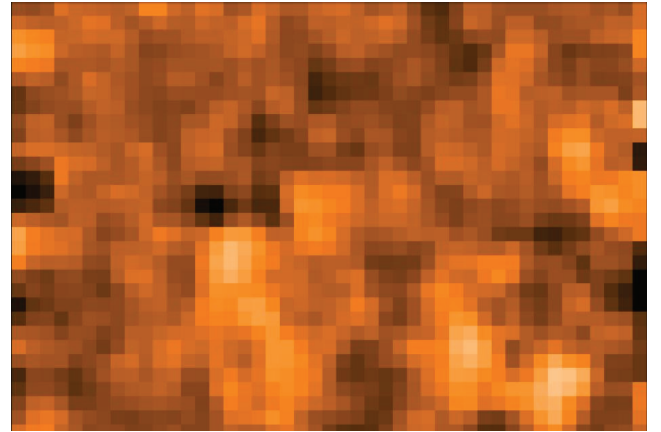


Figure 3. Co-added 70- μm data, showing the flux range 5–95 per cent. The averaged star signal is visible at the centre.

bright somewhat longwards,¹ at 70 μm . Since stars observed with good signal-to-noise ratio tend not to be underbright, an optimal scaling test was carried out to check the observed to predicted flux ratio. The scale factor is found to be 1.08 ± 0.06 , so we conclude that there is no general 70- μm flux deficit and the faint results are by chance. Also, the distribution of significance values χ_{70} (Table 2) is a slightly broadened Gaussian ($\sigma = 1.15$) centred close to zero, indicating the 70- μm results are not anomalous.

4.2 Dust excesses

To search more deeply for cool excess among the ensemble, frames for the 11 new 70- μm observations were stacked about the stellar positions, taking the resulting mean per pixel signals. In Fig. 3, the mean photosphere is now faintly visible at centre, although background structures from individual frames are also present. Aperture photometry gives a mean flux of 4.6 ± 1.2 mJy, consistent with minimal debris given the mean 70- μm photospheric flux of 5.2 mJy.

A better estimate of the net cool dust within the whole thick disc sample can be found from a weighted average of the 70- μm

¹ The stars HD 65583, 68017, 144579 and 157214 have also been observed with *Spitzer* IRS and have no excesses out to 34 μm (Lawler et al. 2009).

excesses (Table 2). We define $E70 = R70 - 1$ and weight the values by $1/\text{error}^2$, where the error in $E70$ can be written as $E70/\chi^2$. The net excess for the 19 stars with 70- μm data is 0.07 ± 0.34 , i.e. $\lesssim 40$ per cent above the photosphere at the 1σ upper bound. This result is however dominated by five stars with low-noise data. For a stellar effective temperature of around 5600 K and dust peaking in brightness at 70 μm , i.e. of temperature around 50 K, an excess of <40 per cent gives an upper limit on fractional dust luminosity $L_{\text{dust}}/L_{\text{star}}$ of 4×10^{-6} (Beichman et al. 2006).

Approximately 80 per cent of nearby Sun-like stars are also dust free at this level individually (Greaves & Wyatt 2010). Hence, if the thin and thick disc populations have similar debris incidence and levels, the null result here for only 19 stars is not surprising. In more detail, the expected number of debris detections at 70 μm can be estimated from Fig. 1, taking into account the limiting depths of the thick disc data. The number of detection per 50 per cent wide excess bin is then the fraction in the general (T+B) surveys, scaled to 19 stars, and multiplied by the fraction of thick disc targets where such an excess was detectable. This yields a cumulative 2.1 expected detections, with a Poisson error of ± 1.5 , so finding no debris systems is within the errors. Similarly, there is a 5 per cent incidence of 24- μm excess for Sun-like stars in general (Koerner et al. 2009), so finding none among 22 stars here is consistent with similar populations. Studying trends by age, Gáspár et al. (2009) found a 24- μm debris incidence of approximately 4.1 ± 1.5 per cent at 3–10 Gyr. In theoretical models, some slow fading is predicted, with fractional dust luminosity decreasing e.g. by 25 per cent by 11 Gyr compared to typical mid-age Galactic systems at about 5 Gyr (Löhne, Krivov & Rodmann 2008).

Statistics are ultimately limited by the small numbers of thick disc stars actually passing through the solar vicinity at the present time, and in our volume-limited sample we have already searched for debris among the nearest 15, or 70 per cent, of objects catalogued by Karataş et al. (2005). Fainter levels of debris or cooler dust may be accessible to the *Herschel* satellite, now observing at wavelengths of 70–500 μm .

5 DISCUSSION

The debris incidence in the thick disc is either low or normal; we can rule out a particularly *high* incidence of excess in these metal-poor systems. Recent theoretical work by Johansen, Youdin & Mac Low (2009) has suggested that low-metal systems could have late planetesimal formation as the gas disc is dispersed, in which case more collisional debris might have been expected around the thick disc stars, but this is not seen here. In contrast, some thick disc stars do host giant planets – an outcome thought to reflect more efficient planetesimal building at early stages. In our sample, HD 114729 has a gas giant (Butler et al. 2003), and four other good planet candidates exist for thick disc stars (Gonzalez 2009). The M_{Jup} of these planets range from 0.5 to 7 Jupiter masses and the semimajor axes from 0.15 to 2 au.

Fig. 4 shows the predictions of our model (Greaves et al. 2007) based on the mass of solids available at the planet(esimal) formation stage. The model is based on a threshold of solid material needed for a successful outcome such as building a gas giant core, and with the solid mass within the disc scaling by the metal fraction in common with the star. The M_{solids} distribution is obtained by multiplying the lognormal distribution of overall disc masses (Andrews & Williams 2005) by the lognormal distribution $[\text{Fe}/\text{H}]$ for the population of interest, here with mean and standard deviation of -0.6 and 0.2 , respectively. The thresholds for forming detectable debris discs and

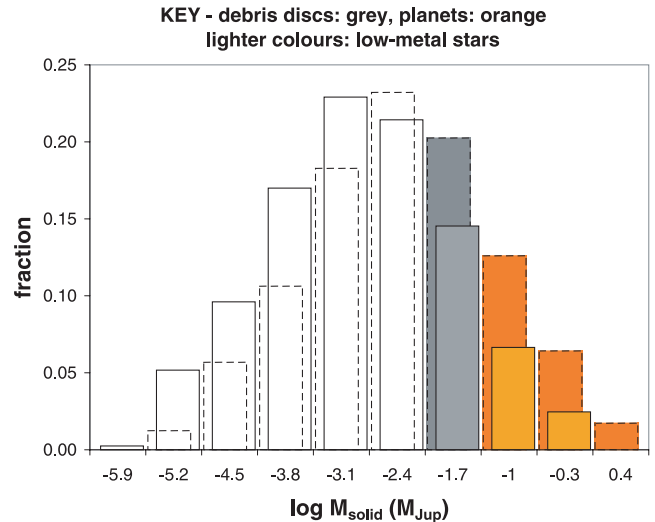


Figure 4. Distributions of solid masses available in the discs at early times (see text), for thin disc stars and the general nearby star population (shown with dashed boxes and darker shading). Giant planets can form from the three rightmost bins (orange) and debris discs from the adjacent bin (grey); mass labels refer to the lowest values in the bin.

gas giants are around 0.02 and 0.1 Jupiter masses in solids, respectively (Greaves et al. 2007). For simplicity we neglect systems with both planets and debris, since the two phenomena appear uncorrelated (Bryden et al. 2009), and so this is a small subset of Sun-like stars.

Fig. 4 then predicts that 15 per cent of thick disc stars should form giant planets, with a further 9 per cent forming comet belts. Around the general local population of Sun-like stars, the corresponding numbers are 21 and 20 per cent. The differences are modest because the availability of solid material is largely dictated by the total disc masses, which have a very broad distribution of $\sigma = 1.15$ dex, compared to the standard deviations around 0.2 dex for the metallicities. Hence, while the peaks of the solid-mass distributions do shift for differing metals (Fig. 4), the upper tails associated with forming planet(esimals) are broad and do not have greatly different numbers of stars.

These values are compatible with survey data, taking into account that not all systems of planets and debris have yet been discovered. In the general stellar population, the extrapolated planet total is 17–20 per cent (Cumming, Marcy & Butler 2009) and the debris disc count at 70 μm is around 14–20 per cent (Trilling et al. 2008; Greaves & Wyatt 2010), which agree well with the predictions of 21 and 20 per cent, respectively. For the thick disc stars, the debris incidence is $\lesssim 5$ per cent versus the 9 per cent predicted, and the planet count is estimated at ~ 6 –11 per cent, versus 15 per cent predicted. It has been noted that lower-metal discs may disperse faster (Yasui et al. 2009), and this could reduce the efficiency of growing planet(esimals) compared to our simple model.

The thick disc planet frequency is poorly known, with our 6 per cent estimate derived from three planetary systems among 52 thick disc objects that have $[\text{Fe}/\text{H}] < -0.2$ and $[\alpha/\text{Fe}] > 0.2$ (Valenti & Fischer 2005; Reid et al. 2007), while the 11 per cent is for two out of 18 systems in the Fe-, Mg-selected, volume-limited, thick disc sample of Fuhrmann & Bernkopf (2008). These planet frequencies are actually *higher* than for low-metal stars in general, such as the incidence of 1 per cent for periods of less than 3 yr found in a recent survey by Sozzetti et al. (2009). Conditions that

could promote the formation of more gas giants in the old thick disc population compared to younger low-metal thin disc stars are at present unclear.

6 CONCLUSIONS

While the scarcity of nearby thick disc stars makes it difficult to study the population, there are examples of disc and planet systems around some comparably old stars. A gas giant is known around the post-main-sequence star V391 Peg of age $\gtrsim 10$ Gyr (Silvotti et al. 2007); the ≈ 10 -Gyr old G8 V star τ Ceti has a debris disc (Greaves et al. 2004) and the oldest known star with both a planet and debris is 70 Vir at ≈ 9 Gyr (Bryden et al. 2009). These discoveries point to some of the first stars in the Galaxy having had the potential to form planetary systems akin to those around typical stars seen today. If some thick disc stars have minimal cool debris but do host gas giants orbiting at a few au, these would be analogues to the Solar system formed many Gyr before the Sun.

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